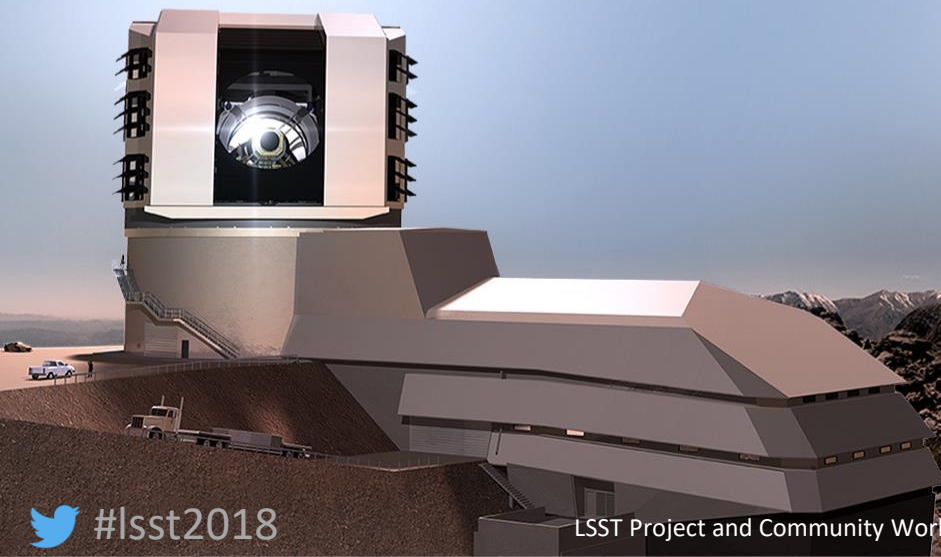




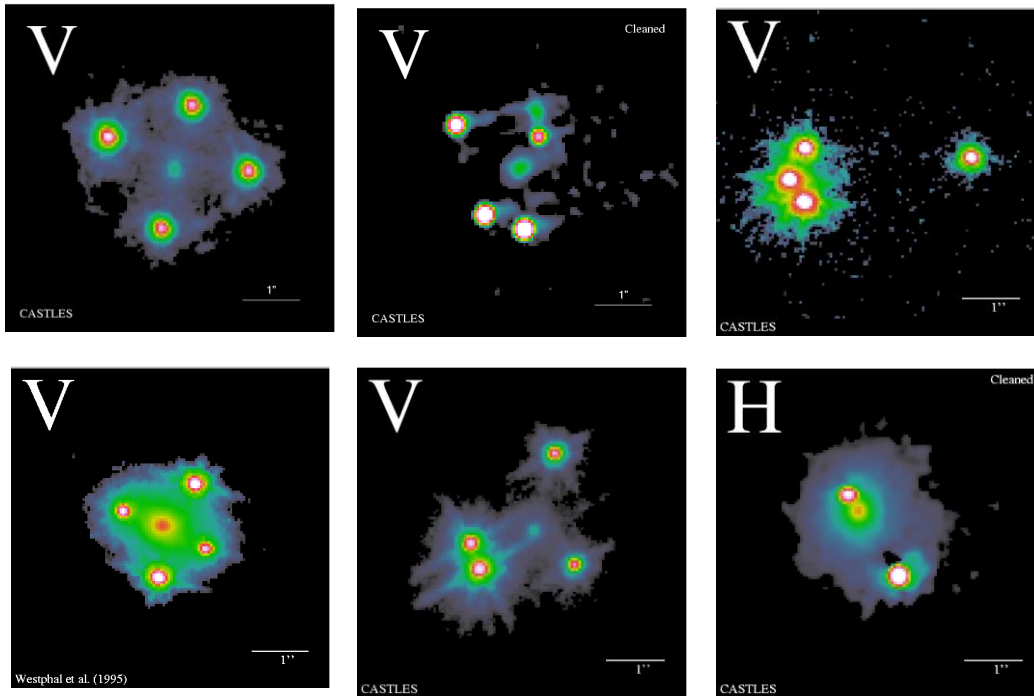
Supermassive BH Science: LSST Science Expectations and Cadence Needs for Strongly Lensed AGN (or quasar microlensing with the LSST)

Timo Anguita
UNAB/MAS, Chile

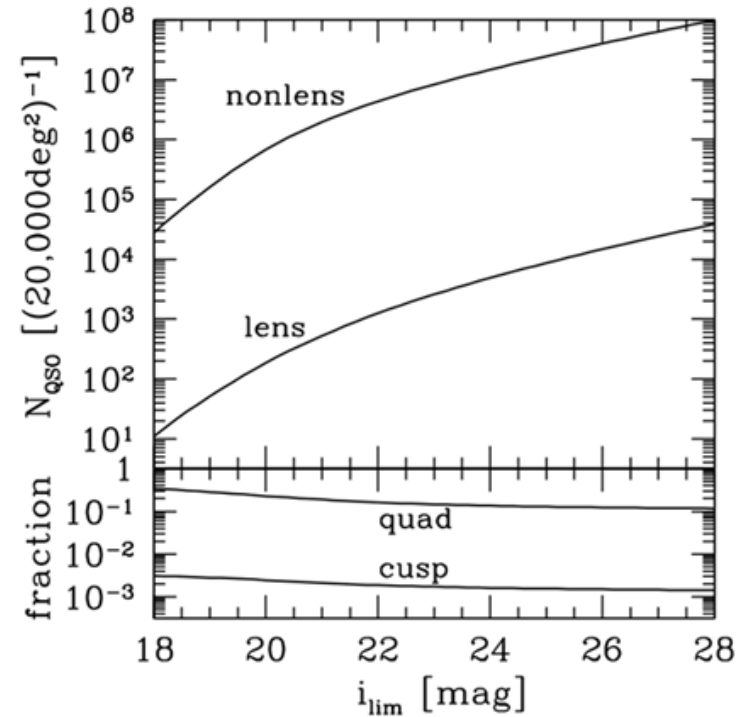
August 16, 2018



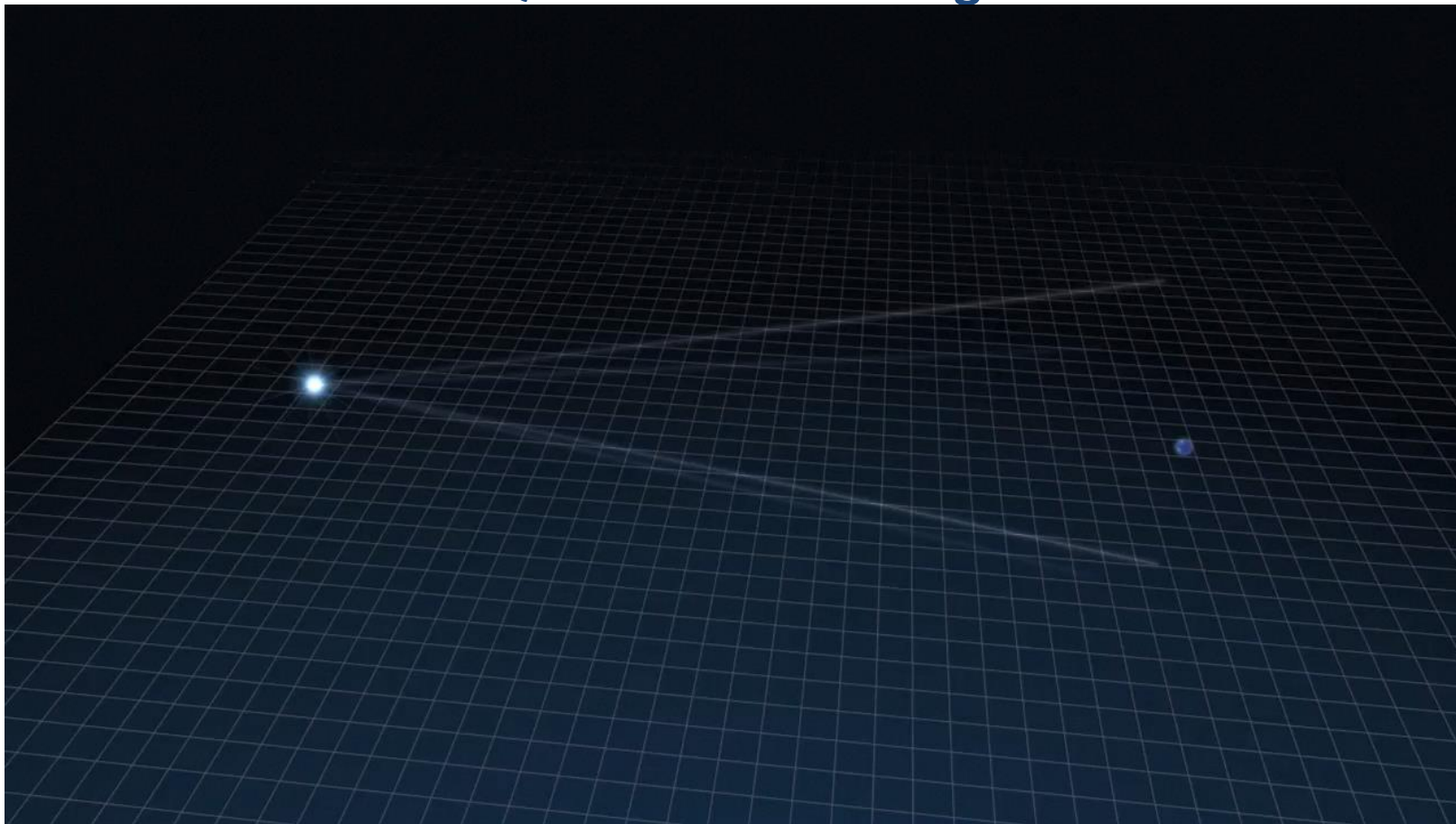
Lensed Quasars



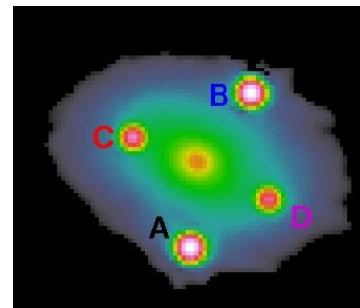
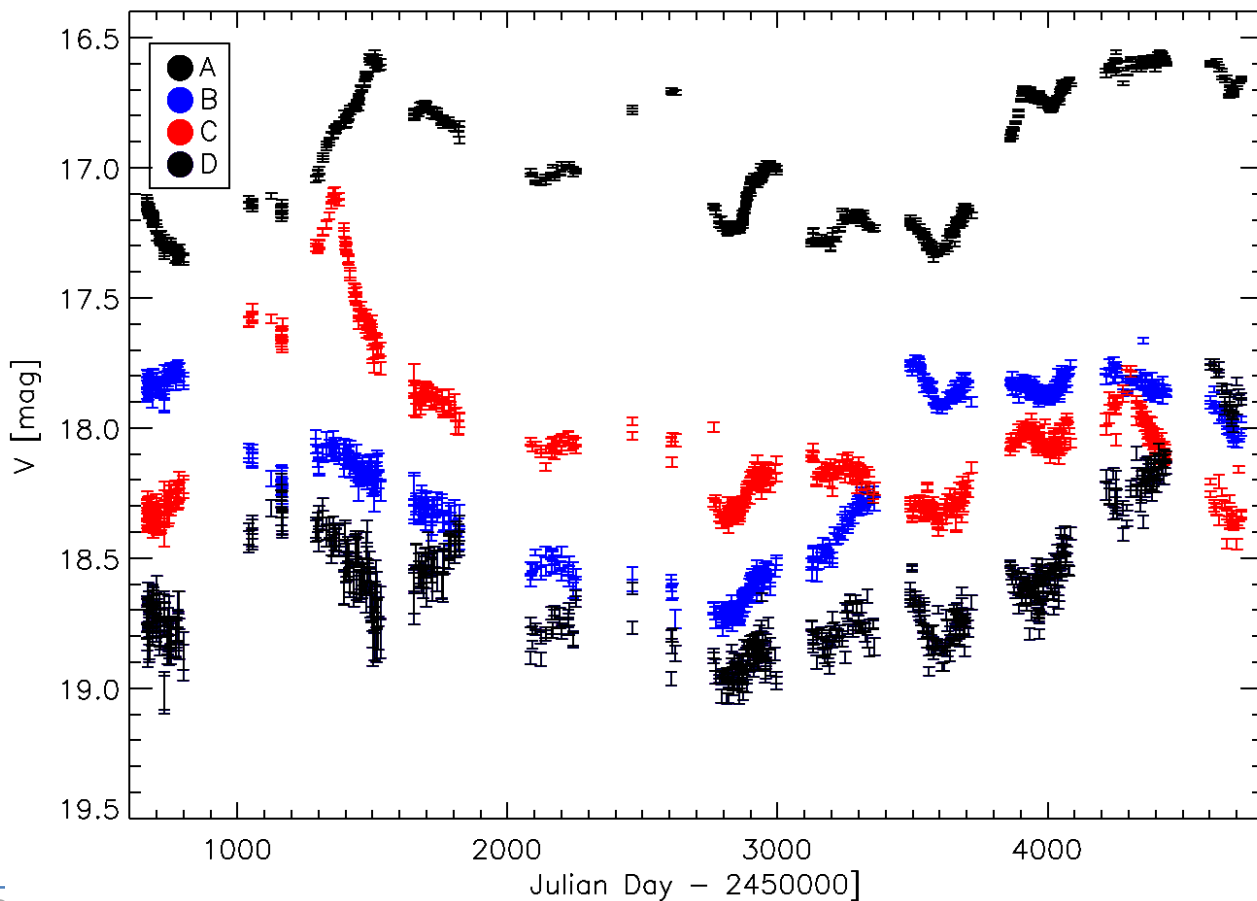
Oguri and Marshall (2010)



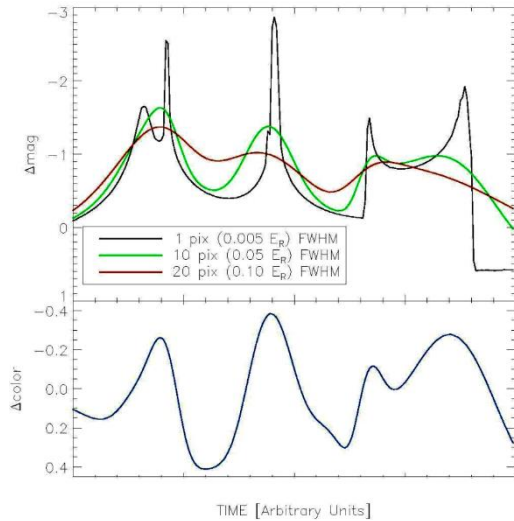
Quasar Microlensing



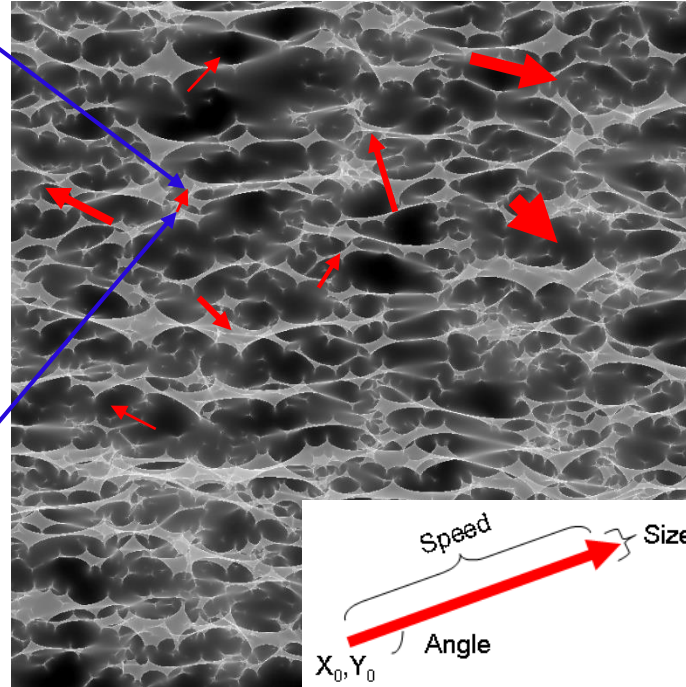
Q2237+0305 (10 years)



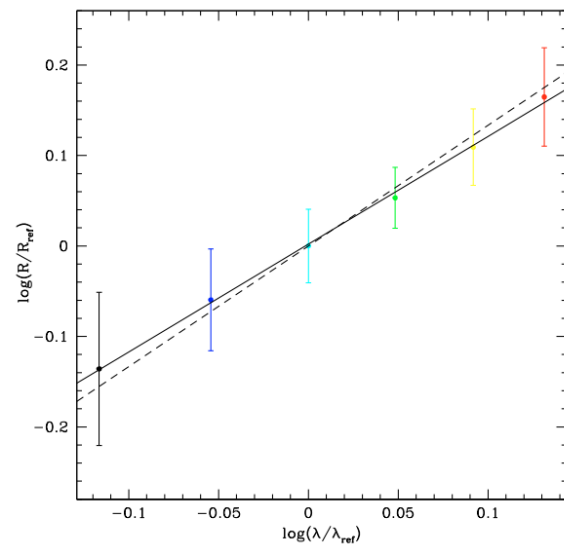
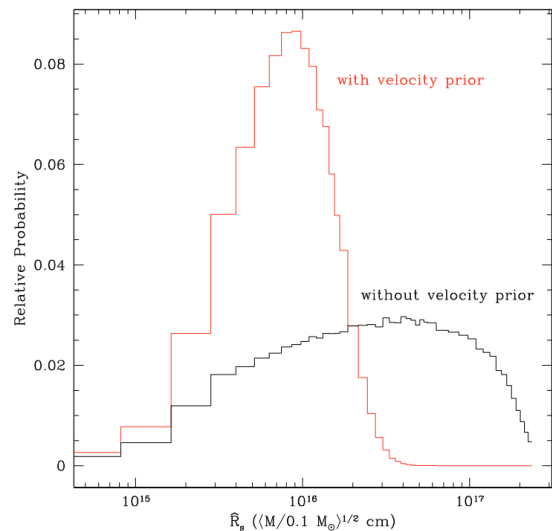
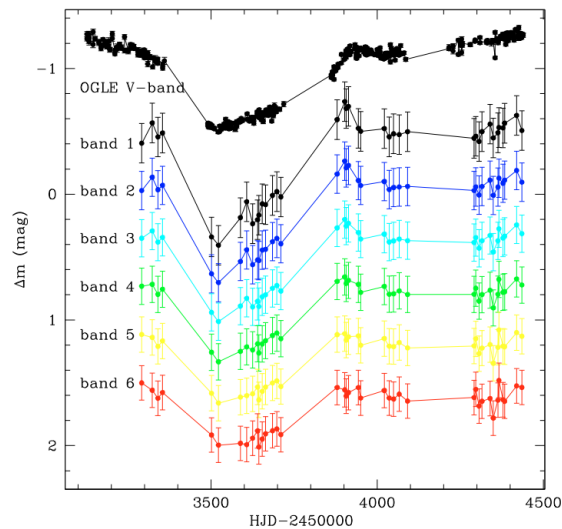
Microlensing Pattern Analysis



- Different source sizes show different light curves
- Difference between two shows color curve

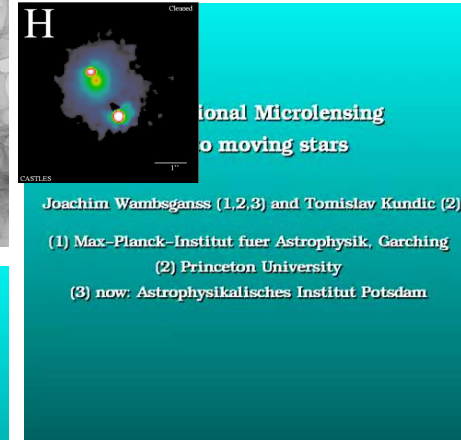
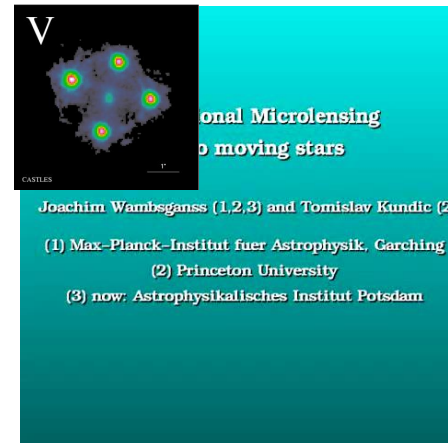
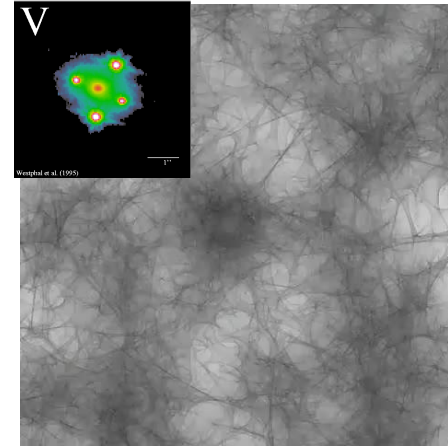


Example: Multi-band Single High Magnification Event (Eigenbrod+2008)

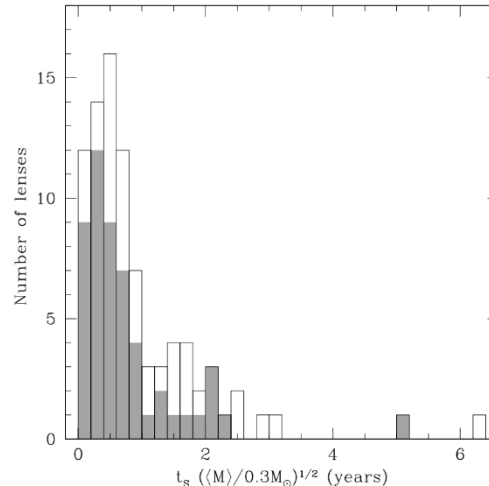
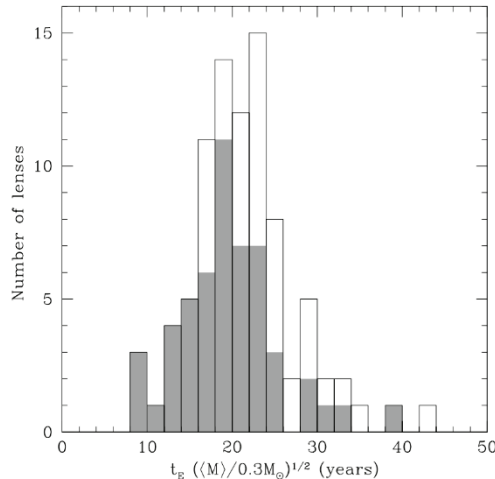
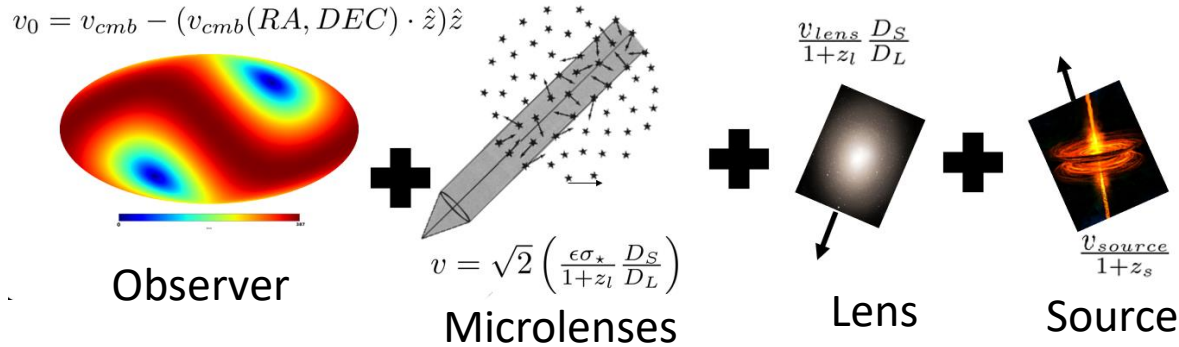


Number of microlenses

- Strong mass model in each image
 - κ , surface mass density \rightarrow amount of mass
 - γ , shear \rightarrow “gravitational pull”
- Further follow-up
 - s , mass in DM
 - $\kappa^*(1-s) \rightarrow$ amount of stars



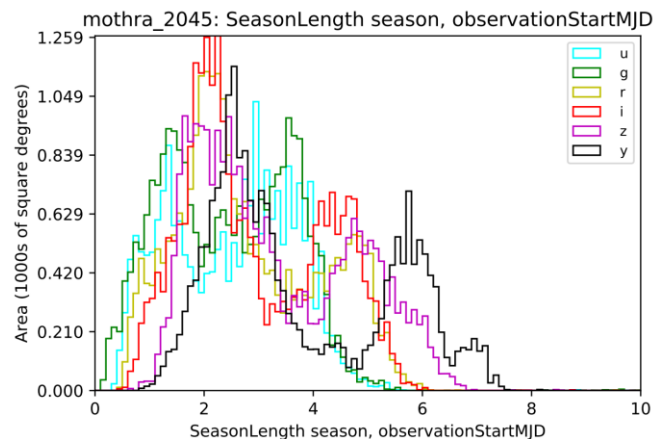
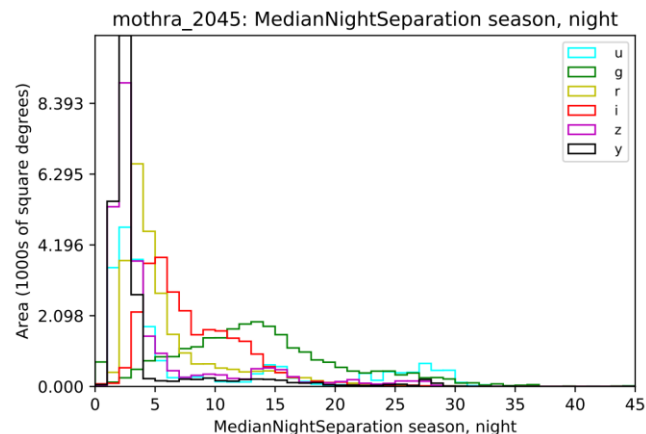
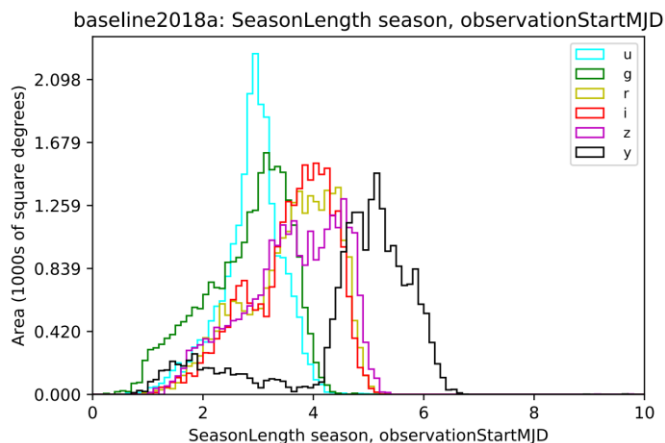
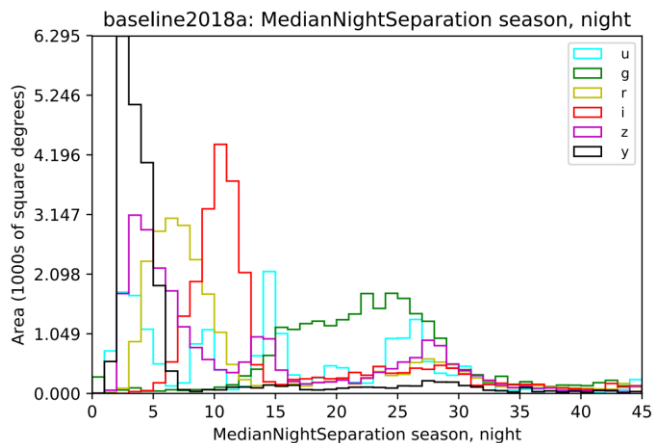
Velocities and Time-Scales



- Mosquera et al. 2011, known systems
 - $t_E \rightarrow$ Time required to cross $E_R \rightarrow$ frequency of events
 - $t_s \rightarrow$ Time for a “caustic” to cross a i-band-sized source \rightarrow length of an event

High level differences (MAF analysis)

Baseline

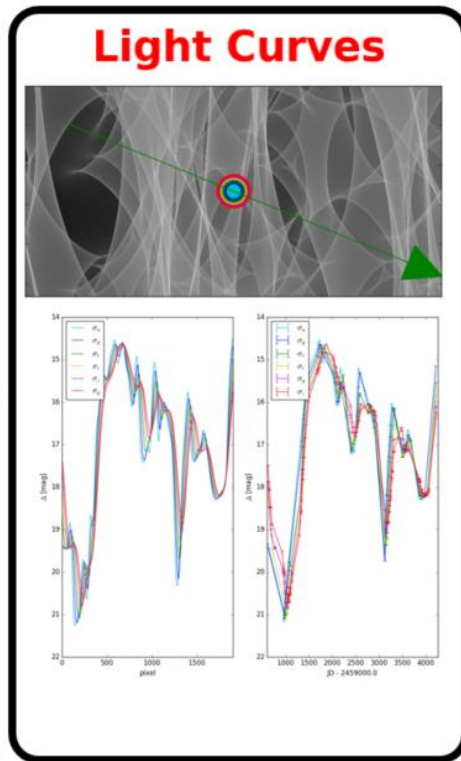
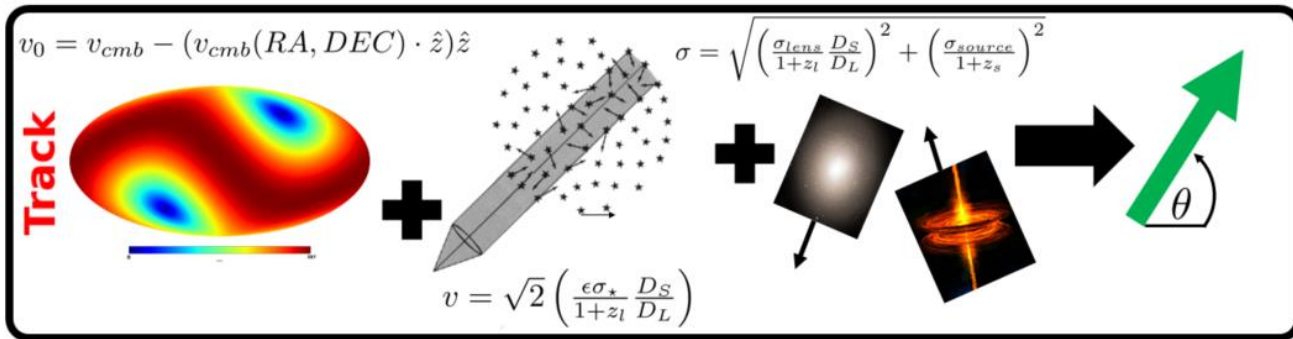
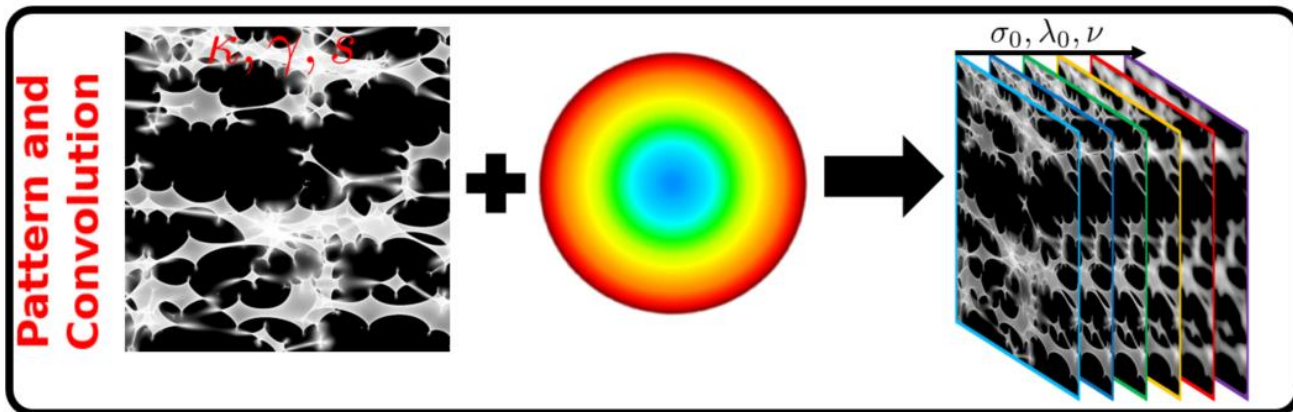


Rolling



The LSST Quasar Microlensing Simulator

RECIPE

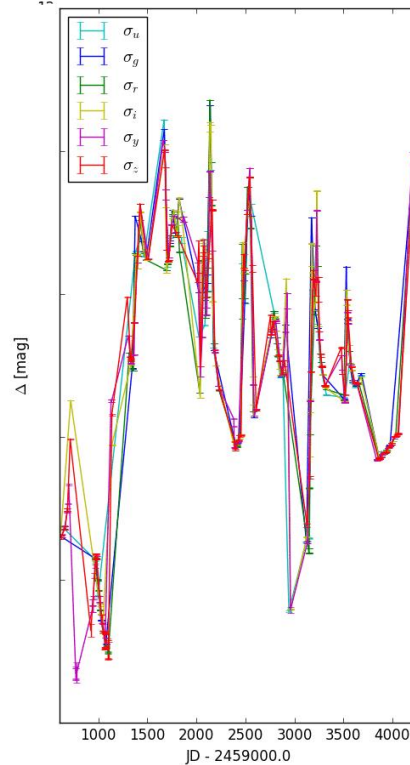
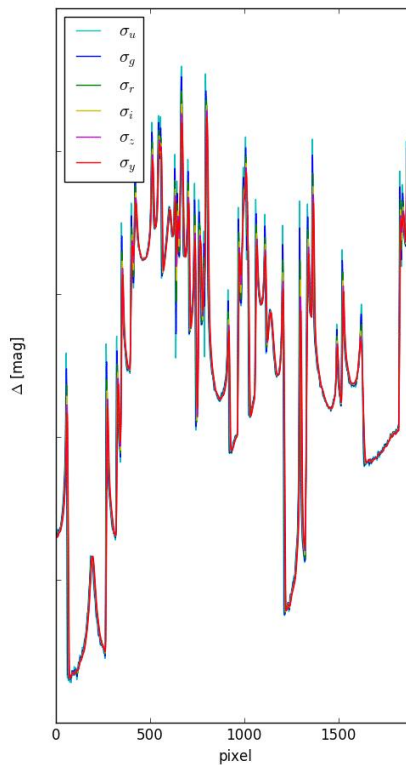
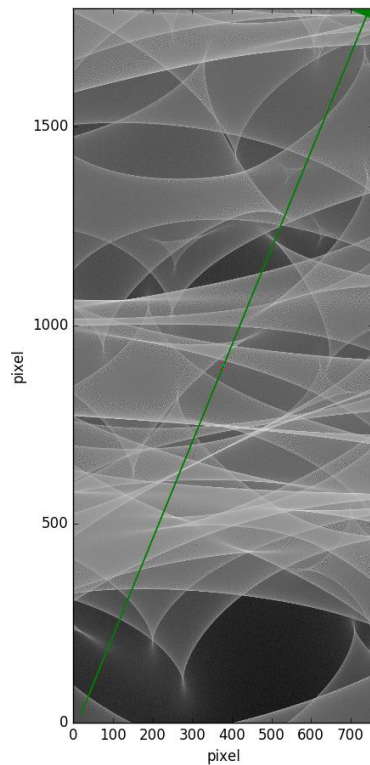


Neira, Anguita & Vernardos 2018 (in prep)

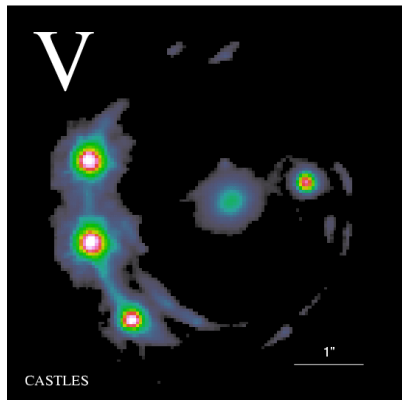


The LSST Quasar Microlensing Simulator

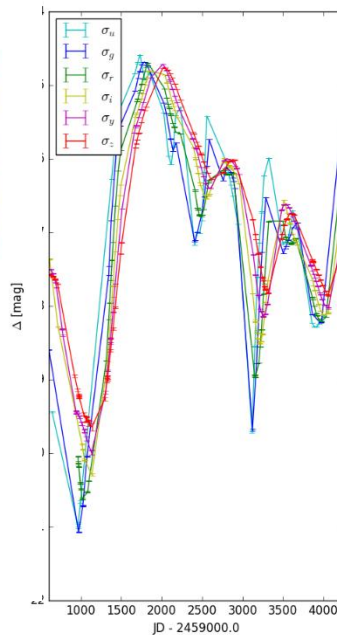
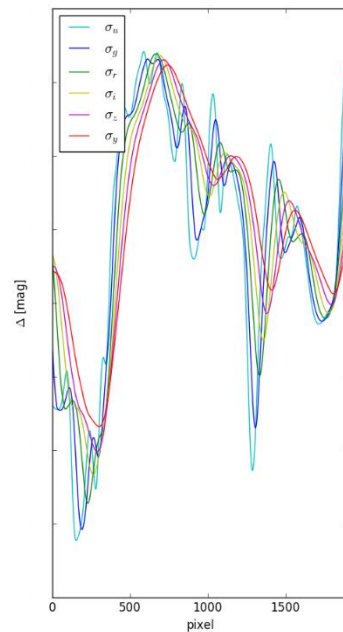
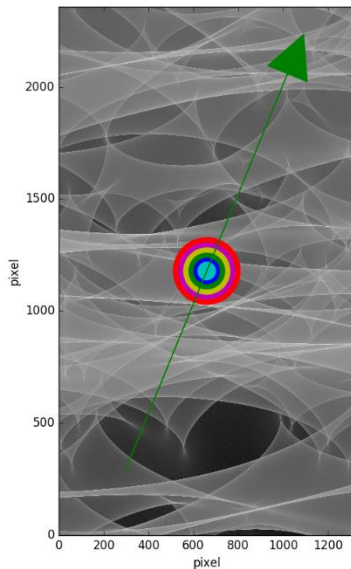
$\nu=0.9$ $\sigma_0=0.1[\text{ld}]$ @ 1026.8\AA



RXJ1131-1231



$l=1.33 \quad \sigma_0=1.0[|d|] @ 1026.8\text{\AA}$



$$v_t \tilde{\propto} D_S/D_L = 1.6x$$

$$\kappa = 0.494$$

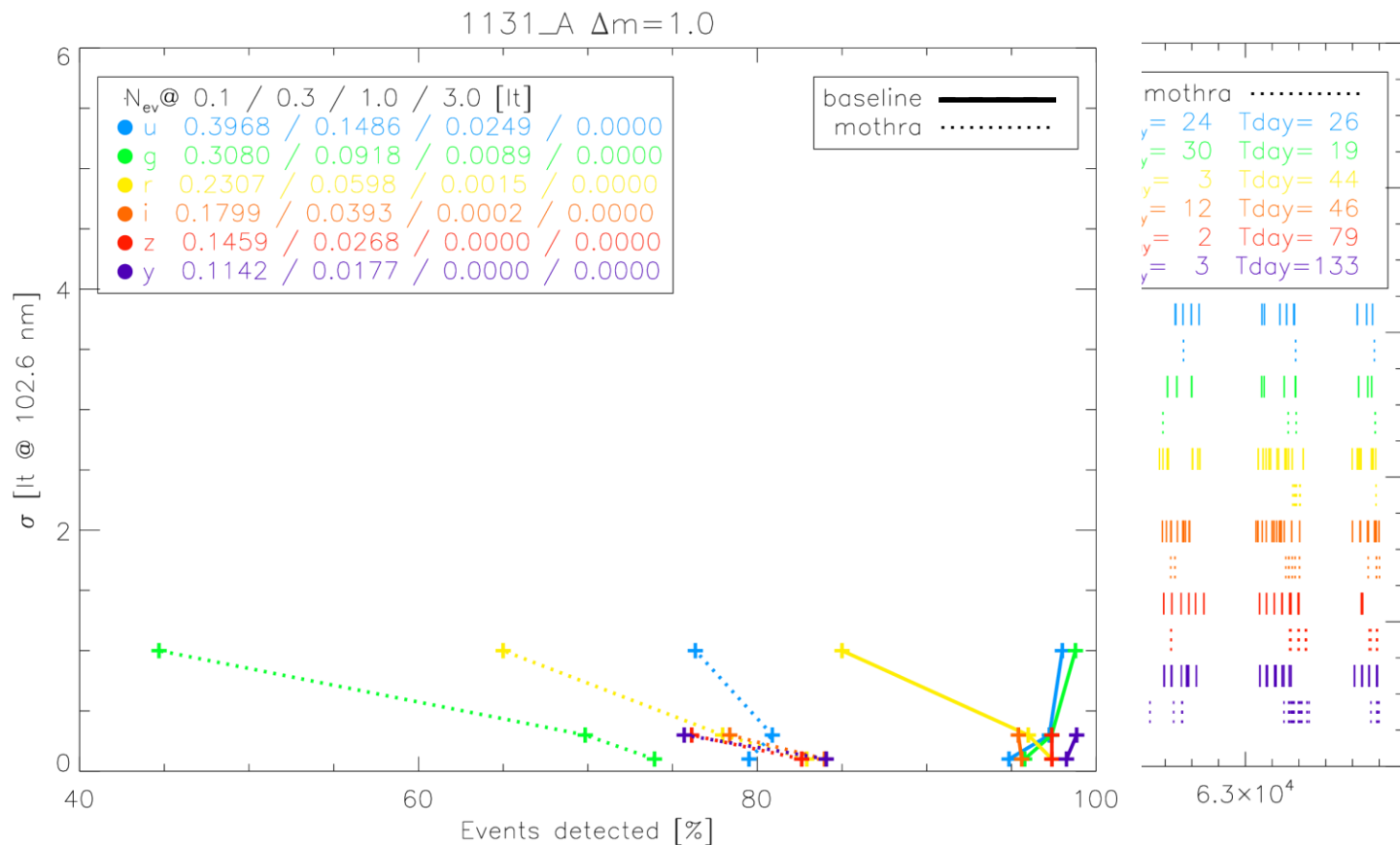
$$\gamma = 0.562$$

$$s = 0.6$$

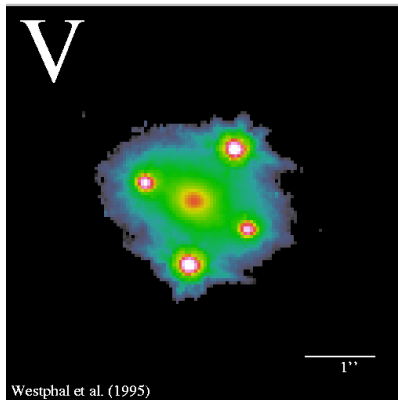
$$\kappa [\text{amount of mass}] * (1 - s [\% \text{ of mass in DM}]) = 0.2x \propto N_{\text{micro}}$$



RXJ 1131-1231



Q2237+0305



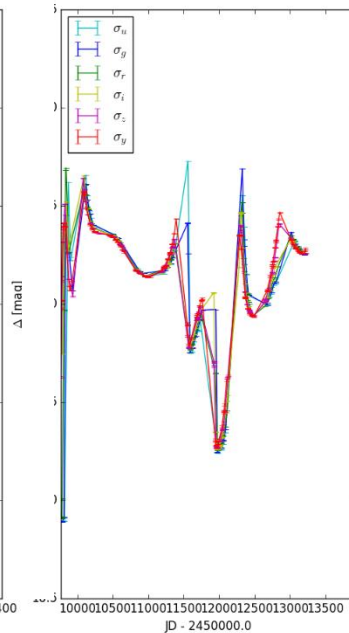
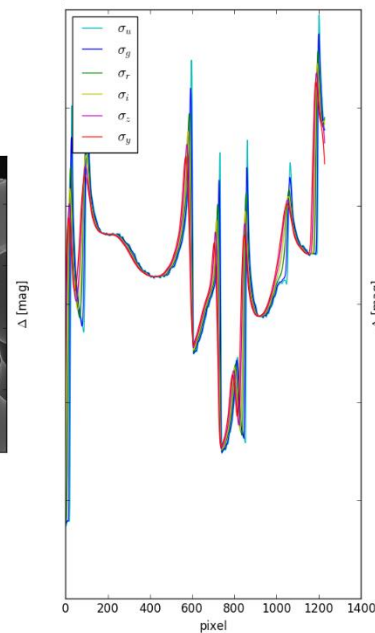
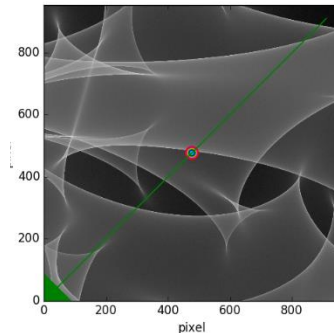
$$v_t \tilde{\alpha} D_S/D_L = 12.5x$$

$$\kappa = 0.394$$

$$\gamma = 0.395$$

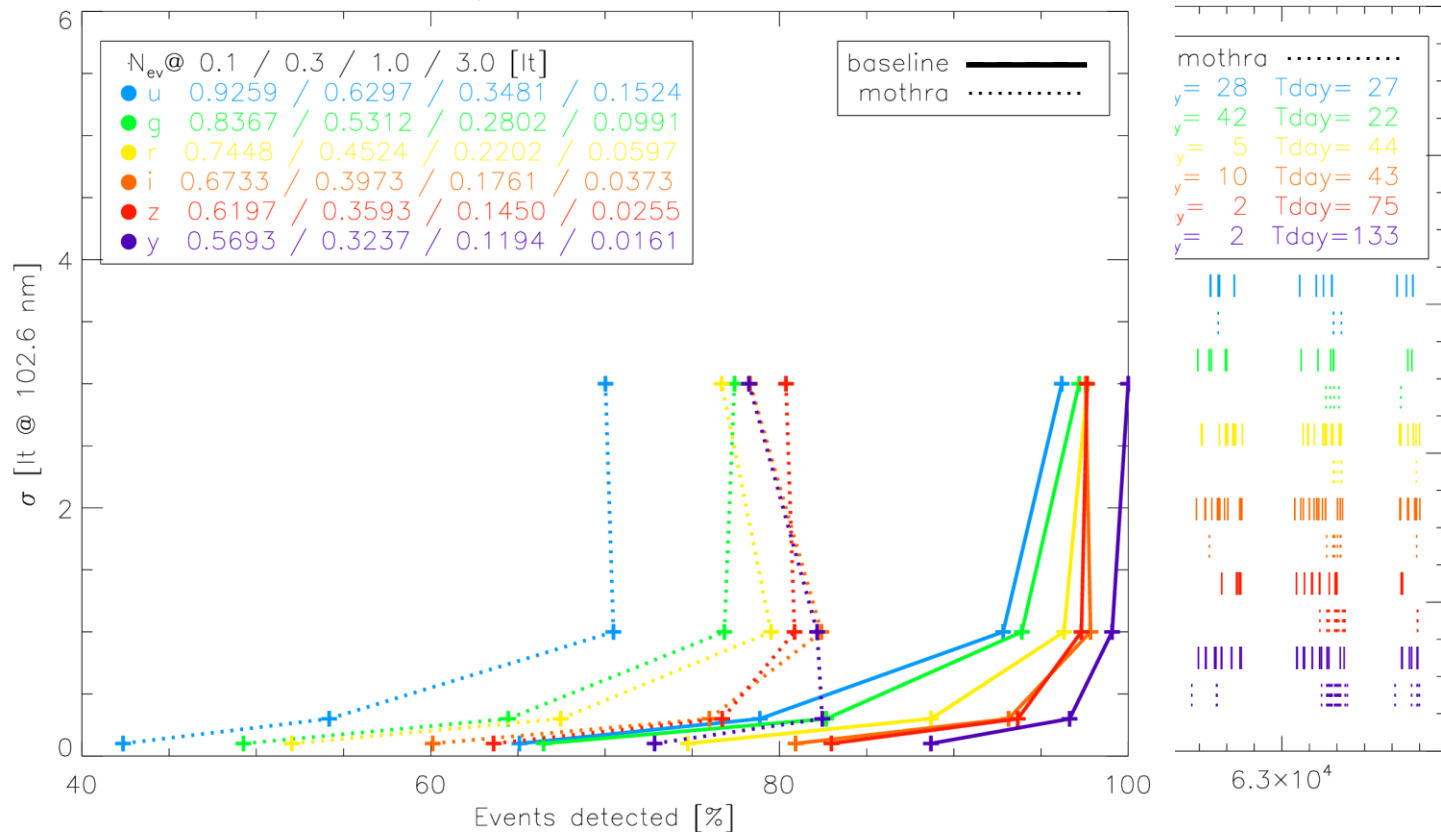
$$s = 0.0$$

$$N_{\text{micro}} \propto 0.394x$$



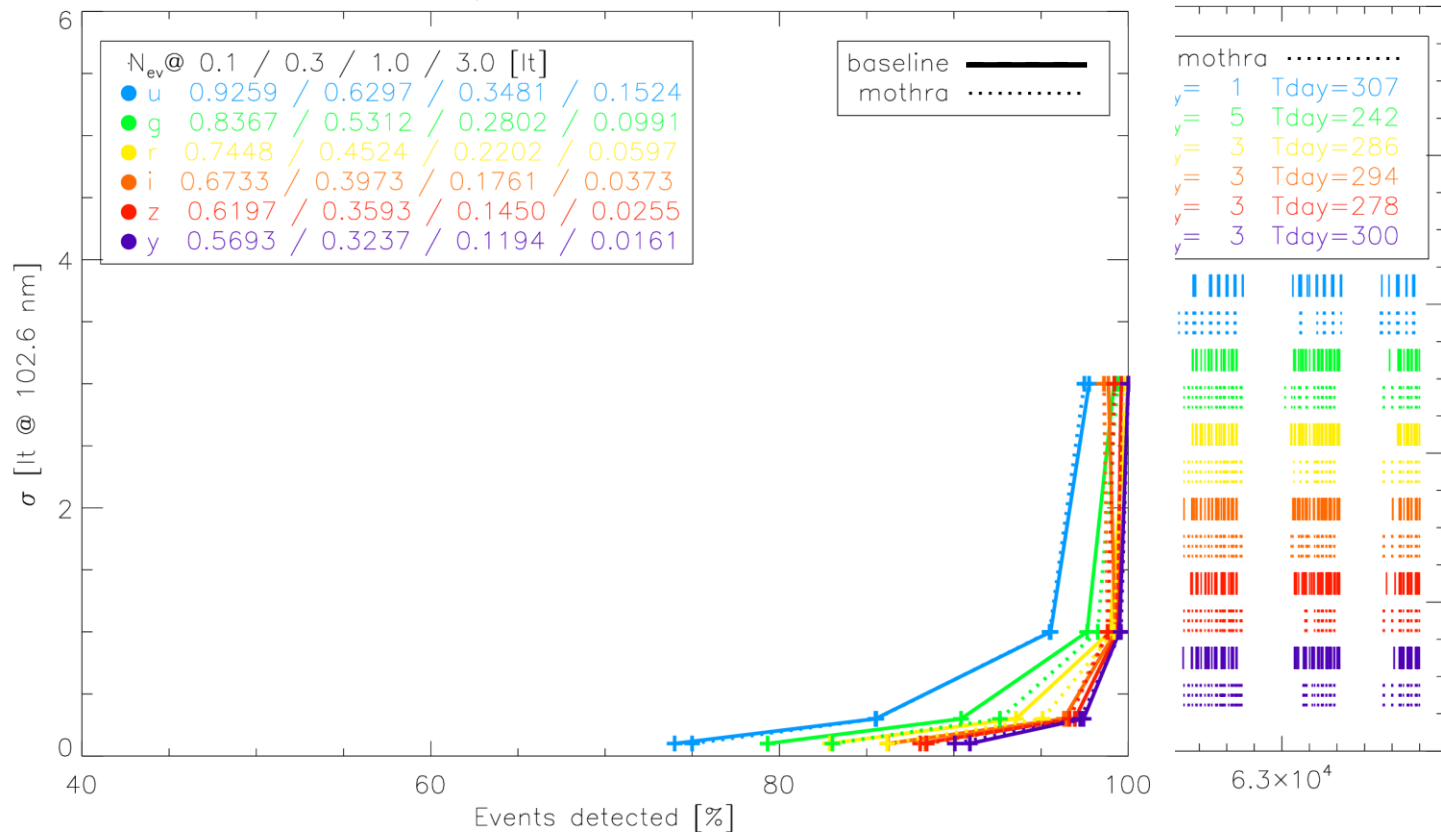
Q2237+0305

Q2237_A $\Delta m=1.0$



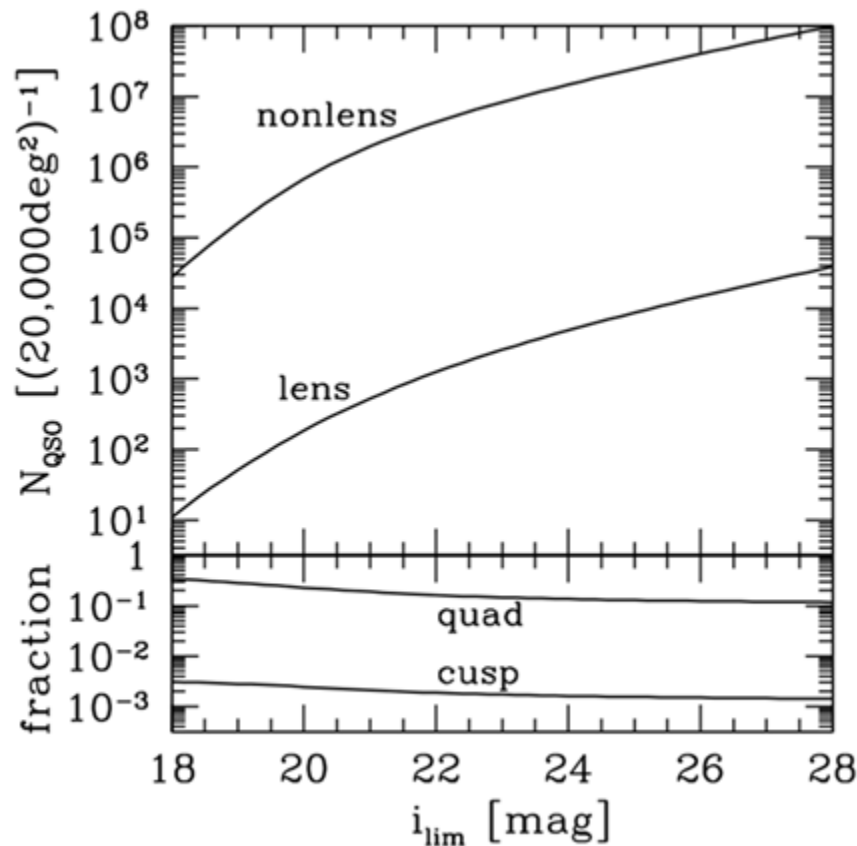
Q2237+0305 in a DDF

Q2237_B $\Delta m = 1.0$



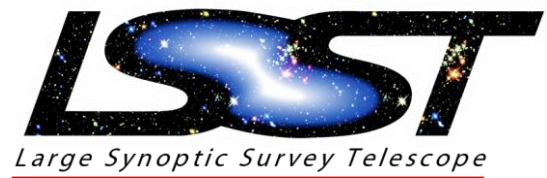
LSST: How many microlensing “usable” lensed quasars?

- All will have microlensing → maximize number of (bright systems)
 - Good seeing g images maximize discovery (Collet 2015)
 - Increase WFD area?
 - Minimize galactic center (in u and g)?
- Oguri and Marshall 2010
 - 8000 lensed quasars (15% quads)
 - 3000 with measured time delays
- Less uniform cadence → same number of systems, perhaps smaller fraction of time delays measured? (not entirely clear still)

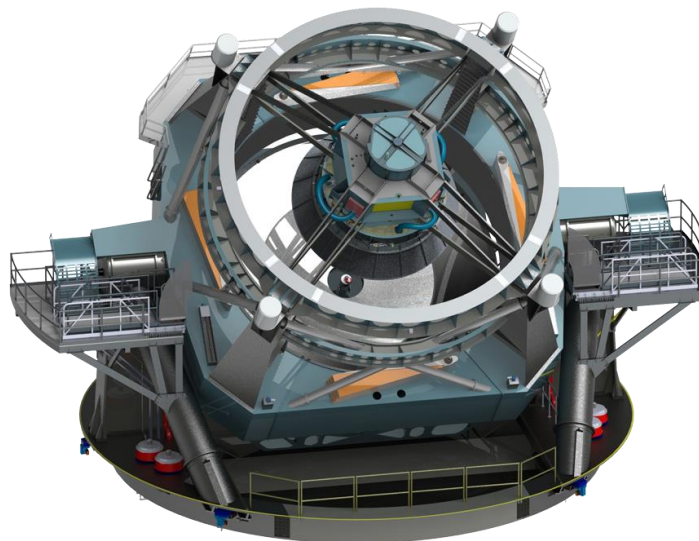


Summary

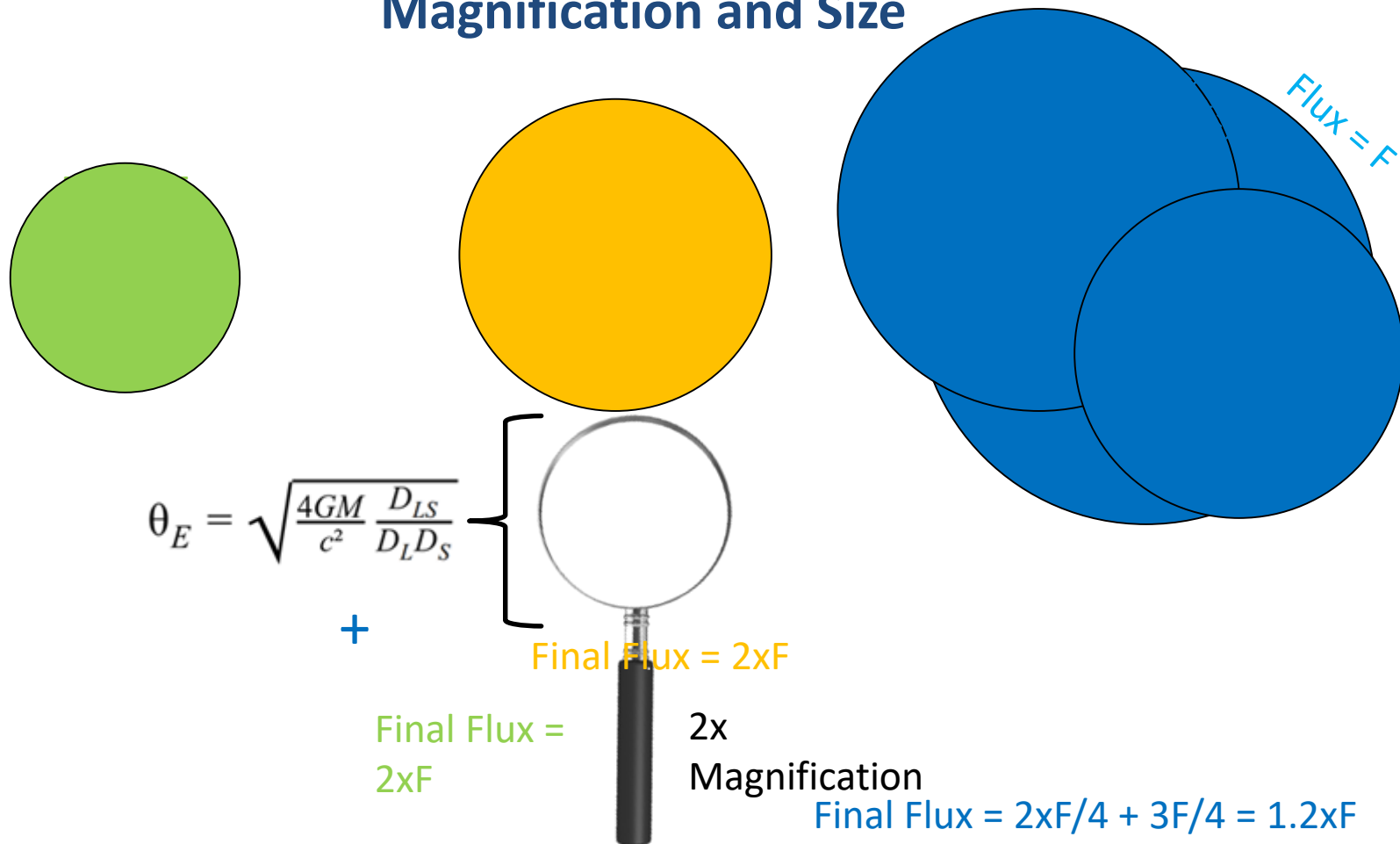
- LSST will be able to produce light curves with significant microlensing signal for thousands of quasars
- Scarcity of systems, long timescales → Ideal survey
- A rolling cadence
 - Less uniform + short seasons → Loss of some “chromatic” events, loss of bluer bands cadence
 - would reduce the microlensing signal
- But also, time delay measurements are necessary for (most) microlensing studies.



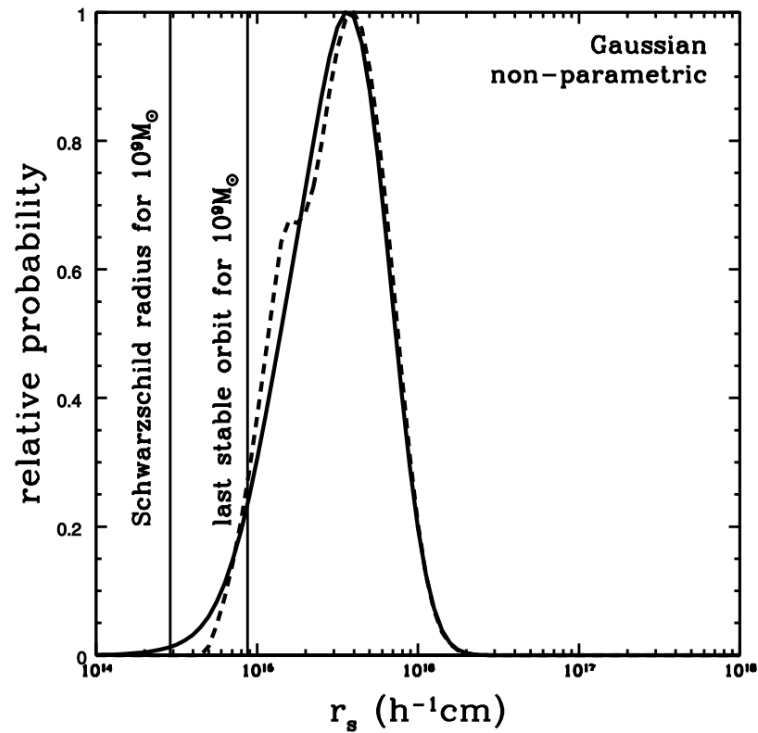
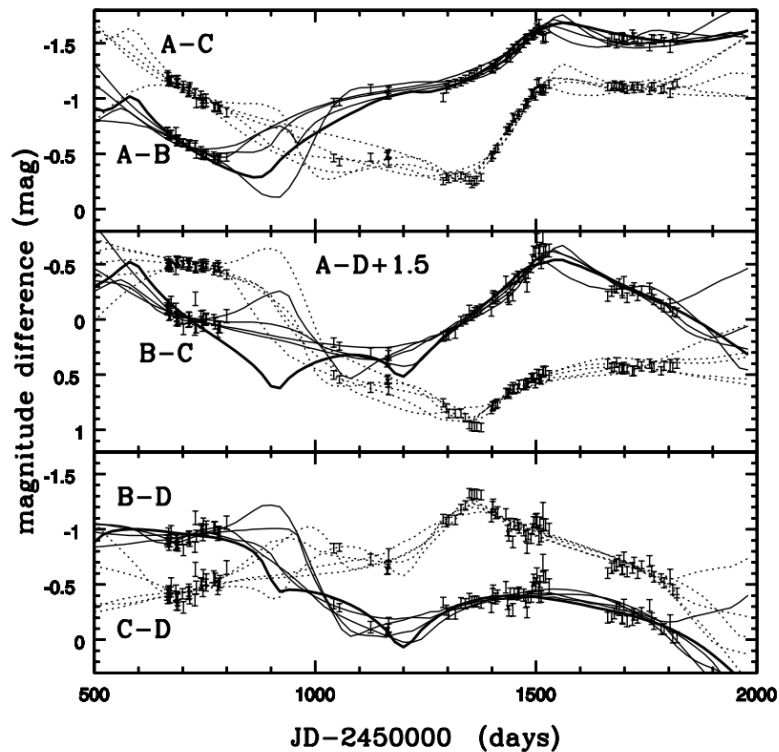
Thanks!



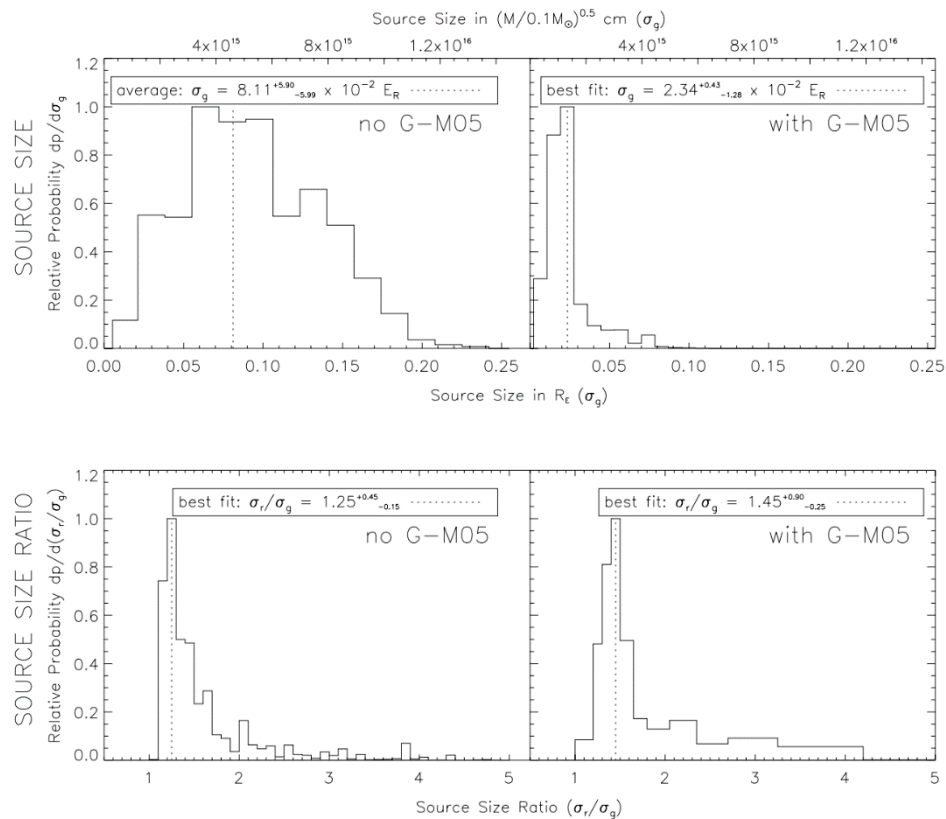
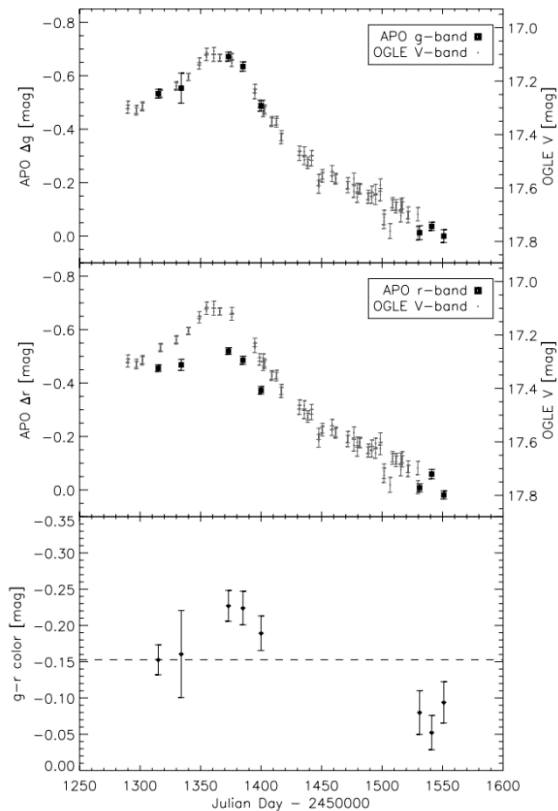
Magnification and Size



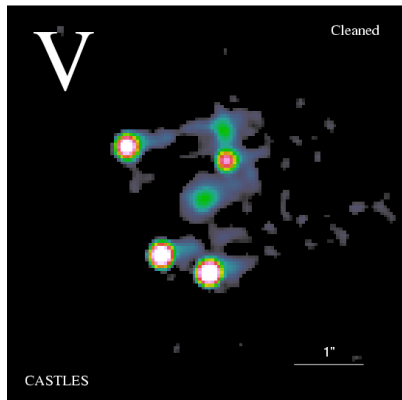
Example: Long Light Curves (Kochanek 2004)



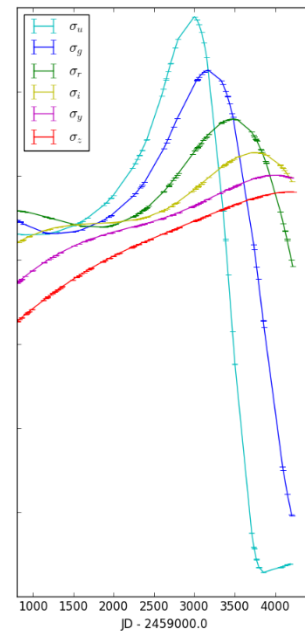
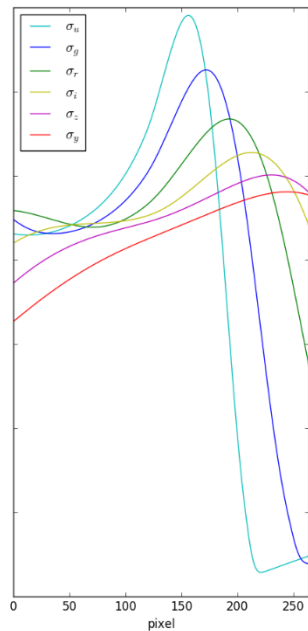
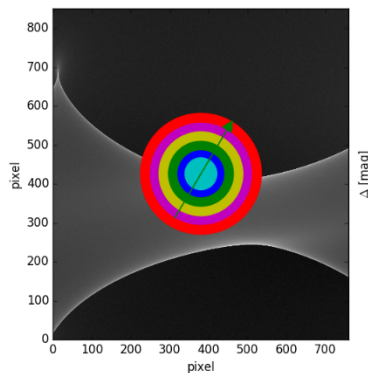
Example: Two band Single High Magnification Event (Anguita+2008)



HE0230-2130



$\nu=1.33$ $\sigma_0=1.0[|d|]$ @ 1026.8Å



$$v_t \tilde{\propto} D_S/D_L = 1.3$$

$$\kappa = 0.472$$

$$\gamma = 0.416$$

$$s = 0.99$$

$$N_{\text{micro}} \propto 0.005$$



HE 0230-2130

